RADIO-CONTINUUM EMISSION FROM THE YOUNG GALACTIC SUPERNOVA REMNANT G1.9+0.3

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SUMMARY: We present an analysis of a new Australia Telescope Compact Array (ATCA) radio-continuum observation of supernova remnant (SNR) G1.9+0.3, which at an age of \( \sim 181 \pm 25 \) years is the youngest known in the Galaxy. We analysed all available radio-continuum observations at 6-cm from the ATCA and Very Large Array. Using this data we estimate an expansion rate for G1.9+0.3 of \( 0.563 \pm 0.078 \)% per year between 1984 and 2009. We note that in the 1980’s G1.9+0.3 expanded somewhat slower (0.484% per year) than more recently (0.641% per year). We estimate that the average spectral index between 20-cm and 6-cm, across the entire SNR is \( \alpha = -0.72 \pm 0.26 \) which is typical for younger SNRs. At 6-cm, we detect an average of 6% fractionally polarised radio emission with a peak of \( 17 \pm 3 \)% . The polarised emission follows the contours of the strongest of X-ray emission. Using the new equipartition formula we estimate a magnetic field strength of \( B \approx 273 \) \( \mu \)G, which to date, is one of the highest magnetic field strength found for any SNR and consistent with G1.9+0.3 being a very young remnant.

Key words. ISM: individual objects: G1.9+0.3 - ISM: supernova remnants - radio continuum: ISM – supernovae: general

1. INTRODUCTION

It is widely accepted that current catalogues have a distinct deficit of young Galactic supernova remnants (SNRs), that is, SNRs \( < 2000 \) years old, with only \( \sim 10 \) confirmed out of a predicted \( \sim 50 \) (van den Bergh and Tammann 1991, Cappellaro 2003). Of these confirmed SNRs, G1.9+0.3 is of particular interest as it is believed to be the youngest in the Milky Way (MW), i.e. \( \sim 150 \) years old (Reynolds et al. 2008, Green et al. 2008, Reynolds et al. 2009, Carlton et al. 2011a,b). Originally identified as a probable SNR by Green and Gull (1984) at 4.9 GHz using the Karl G. Jansky Very Large Array (VLA), G1.9+0.3 was described as a shell source with an approximate brightness slightly less than that of the Tycho and Kepler...
SNRs with a spectral index of $\alpha \sim -0.7^1$. Using the Molonglo Observatory Synthesis Telescope (MOST) Galactic Survey data, Gray (1994) confirmed the classification of G1.9+0.3 as an SNR: the source was described as featuring a shell-like morphology in the radio with an estimated diameter of $1'2$. Later, LaRosa et al. (2000) produced a 90-cm image of G1.9+0.3 using observations from 1985. They estimated the 20/90-cm spectral index of the SNR to be $\alpha = -0.93 \pm 0.25$ and the angular diameter to be $1'1$. Nord et al. (2004) revisited the data collected by LaRosa et al. (2000) and – through the application of superior data reduction techniques – measured the diameter of G1.9+0.3 to be $<1'$ and placed it at a distance of $<7.8$ kpc. Green (2004) estimates the diameter of G1.9+0.3 based on 1.49 GHz VLA observations made in 1985 to be $1/2$. Most recently, Roy and Pal (2014) measured the HI absorption distance using known anomalous velocity features near the Galactic Centre (GC) and found a lower limit on G1.9+0.3 distance from Sun as $10$ kpc, some 2 kpc further away from the GC. Therefore, multiple radio observations have confirmed that G1.9+0.3 has the smallest angular diameter for a known Galactic SNR, indicative of its young age.

Green et al. (2008) re-observed G1.9+0.3 at 4.86 GHz using the VLA after Reynolds et al. (2008) used 2007 Chandra images to show G1.9+0.3 had expanded significantly since 1985 and its X-ray emission appeared to be predominantly synchrotron in nature. By comparing these new VLA observations with the 1985 VLA observations made at 1.49 GHz, Green et al. (2008) determined that G1.9+0.3 had expanded by $15\% \pm 2\%$ over 23 years ($\sim 0.65\%$ per year). Using the same VLA observations from 1985 and 1989, Gómez and Rodríguez (2009) derived an expansion rate of $0.46\% \pm 0.11\%$ and an age of $220 \pm 239$ years. By comparing 2007 and 2009 Chandra X-ray images and utilising a simple uniform-expansion model, Carlton et al. (2011a,b) find an expansion rate of $0.642\% \pm 0.049\%$ yr$^{-1}$ and a flux increase of $1.7\% \pm 1.0\%$ yr$^{-1}$, ageing the remnant at $156 \pm 11$ yr assuming no deceleration. Murphy, Gaensler and Chatterjee (2008) found that G1.9+0.3's flux density at 843 MHz increased by $1.22 \pm 0.24\%$ per year over the last two decades.

Borkowski et al. (2013) suggest that G1.9+0.3 was likely a Type Ia SNe with the shell of its remnant in free expansion with a velocity $\sim 18,000$ km s$^{-1}$. The ejecta shows spatial asymmetry with prominent Fe-group elements in the northern rim. Also, we point out that Abramowski et al. (2014) report no $\gamma$-ray signal from G1.9+0.3 using observations from the H.E.S.S. (High Energy Stereoscopic System) Cherenkov telescope array.

The presence of polarised emission and spatial spectral variations are identified by Farnes (2012), with flatter spectra identified in the NW and SE of the remnant.

In this paper, we present the results of our Australia Telescope Compact Array (ATCA) radio-continuum observations of G1.9+0.3 made at 20, 13 and 6 cm in 2009. A comprehensive expansion study is conducted by comparing the new 6-cm observations with a previously unpublished 6-cm ATCA radio-continuum observation, made in 1993, and three 6-cm VLA observations made in 2008, 1989 and 1984 respectively. We also report on the radio-continuum spectral energy distribution and polarisation properties of this young SNR.

2. OBSERVATIONAL DATA

G1.9+0.3 was observed at 20, 13 and 6 cm wavelengths on four days in 2009 (Project C1952). Two of those days being in January 2009 with the remainder in February 2009. The 20-cm and 13-cm observations, taken on the 2nd and 3rd day, were carried out simultaneously as they make use of a common feed-horn. The 6-cm observations taken on the 1st and 4th day, used four different frequencies (two per day) to improve multi-frequency synthesis (MFS). For this purpose, the MIRIAD (Sault and Killeen 2008) task MFPLAN was used to select the most appropriate frequencies. Over the four days G1.9+0.3 was observed with two separate antenna configurations for a total of 29 independent baselines covering a range of spacings from 31 to 6000 m. See Table 1 for complete observational details.

| Table 1. 2009 ATCA observations of G1.9+0.3. |
|------------------|------------------|------------------|------------------|
| **Date** | **Day 1** | **Day 2** | **Day 3** | **Day 4** |
| Date | 03 Jan | 04 Jan | 06 Feb | 07 Feb |
| ATCA Array | 6C | 6C | EW352 | EW352 |
| Frequency 1 | 4.672 GHz | 1.384 GHz | 1.384 GHz | 4.544 GHz |
| Frequency 2 | 5.440 GHz | 2.368 GHz | 2.368 GHz | 5.184 GHz |
| Bandwidth | 128 MHz | 128 MHz | 128 MHz | 128 MHz |
| Time on source | 332 min | 458 min | 509 min | 1033 min |
| Primary Calibrator | J1934-638 | J1934-638 | J1934-638 | J1934-638 |

$^1$Spectral index defined as $S \propto \nu^\alpha$.
Radio-continuum emission from the young galactic SNR G1.9+0.3

Fig. 1. ATCA 20-cm image of Galactic SNR G1.9+0.3. The blue ellipse in the lower left corner represents the synthesised beam of 10′.9 × 5′.4 at PA=−0°.5. Contours are drawn at 3σ, 5σ, 8σ, 12σ, 17σ, 23σ, 30σ, 38σ, 47σ and 57σ (σ = 0.22 mJy/beam).

Fig. 2. ATCA 13-cm image of Galactic SNR G1.9+0.3. The blue ellipse in the lower left corner represents the synthesised beam of 6′.1 × 2′.9 at PA=−0°.5. Contours are drawn at 3σ, 5σ, 8σ, 12σ, 17σ, 23σ, 30σ, 38σ, 47σ and 57σ (σ = 0.32 mJy/beam).

Fig. 3. ATCA 6-cm image of Galactic SNR G1.9+0.3. The blue ellipse in the lower left corner represents the synthesised beam of 2′.8 × 1′.2 at PA=−0°.5. Contours are drawn at 3σ, 5σ, 8σ, 12σ, 17σ, 23σ, 30σ, 38σ, 47σ and 57σ (σ = 0.07 mJy/beam).

In Figs. 1 through 3 we show our new ATCA 2009 images of G1.9+0.3 at 20-cm, 13-cm and 6-cm respectively. All these images were formed using MFS with uniform weighting and were deconvolved using the MIRIAD (Sault and Killeen 2008) CLEAN and RESTOR tasks, with self calibration being applied to the 6-cm image only. We note that our corresponding ATCA flux density measurements are significantly smaller (~50%) than the VLA estimates of Green et al. (2008). We can attribute this large difference to missing short spacings and poorer uv coverage of the ATCA images.

3. RESULTS

3.1. G1.9+0.3 Expansion and Age

A simple way to determine the age (in years) of a young SNR is to compare two images of the SNR taken at different epochs and to measure the percentage expansion the SNR has undergone over the time between the observations. Since we know that the SNR will have expanded by 100% in the intervening period between the SN explosion and the later observation, we can apply the simple formula,

\[
\text{Age} = \frac{100\%}{ER}
\]

where ER is the percentage expansion the SNR has undergone over the time between the observations (in years).

Ideally, to most accurately determine the expansion rate of an SNR, one should compare images from similar observations i.e., at the same wavelength and having similar, if not identical, uv-coverage, resolution and rms noise.

In the case of G1.9+0.3 the first observations were made in 1984 using the VLA, with the best image produced from the 6 cm data (Fig. 4, bottom left). With this in mind, ATCA and VLA archives...
were searched for all available radio-continuum data of G1.9+0.3 made at this same wavelength of 6 cm.

For the expansion study carried out in this paper, one ATCA and three VLA observations made at 6-cm were found at different epochs from the original 1984 to 2008.

The archival ATCA observation was taken on the 10th June 1993 (Project C034; P.I.: A. Gray). This observation was made in the 6-cm band centred at 4672 and 5440 MHz with a bandwidth (BW) of 128 MHz and the telescope in 6A configuration. Source 1934-638 was used for primary calibration and source 1748-253 was used for secondary calibration. The observations were done in the so-called "snap-shot" mode, totalling ~1 hour of integration spread equally over a 12 hour period. With very few short baselines and poor \( uv \)-coverage, resulting \( rms \) of 0.2 mJy beam\(^{-1}\) is the highest amongst all the analysed observations. Consequently, the data for this observation are of very poor quality and, while the results are still presented here, we exclude measurements from this image in our determination of expansion rates and age. The archival VLA observations were from the 1984 (Project AG0146), 1989 (Project AB0544) and 2008 (project AG0793), see Table 2 for details.

All these observations of G1.9+0.3 were reduced and analysed with the MIRIAD (Sault, Teuben and Wright 1995) and KARMA (Gooch 1995) software packages.

As the resolutions of the produced images varied (due to the various array configuration, observational periods and therefore resultant \( uv \)-coverage), the resolutions of all the images were smoothed/convolved to match the image with the lowest resolution (9\(^\prime\) × 4\(^\prime\) at a PA of 0\(^\circ\)), see Fig. 4. Since the \( rms \) noise of the resultant images also varied (see Table 3), it was decided not to try to determine the expansion by looking at how the shock front moved between epochs, but by looking at how the radially averaged shell profile peak moved from epoch to epoch.

Using RA=17\(^{h}\)48\(^{m}\)45.4, Dec=−27°10′06″ as the centre of G1.9+0.3 we produced normalised shell profiles (Fig. 5), averaged over all angles, for each of the images shown in Fig. 4.

From the shell profiles in Fig. 5 we have determined the expansion rate in arc-seconds per year, percentage expansion per year, the averaged expansion velocity in km s\(^{-1}\) (assuming a distance of 8.5 kpc\(^{3}\)), and age in years. These results are summarised in Tables 4 though 7, respectively.

### Table 2. VLA observations of G1.9+0.3.

<table>
<thead>
<tr>
<th></th>
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<th></th>
<th></th>
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<tbody>
<tr>
<td>1984</td>
<td>26 May</td>
<td>23 June</td>
<td>12 March</td>
</tr>
<tr>
<td>VLA Array</td>
<td>C</td>
<td>BC</td>
<td>C</td>
</tr>
<tr>
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<td>4.89 GHz</td>
<td>4.89 GHz</td>
</tr>
<tr>
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<td>50 MHz</td>
<td>50 MHz</td>
</tr>
<tr>
<td>Time on source</td>
<td>~10 min</td>
<td>~10 min</td>
<td>~30 min</td>
</tr>
<tr>
<td>Primary Calibrator</td>
<td>3C286</td>
<td>3C286</td>
<td>3C286</td>
</tr>
<tr>
<td>Secondary Calibrator</td>
<td>J1832-105</td>
<td>J1751-253</td>
<td>J1751-253</td>
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</table>

### Table 3. G1.9+0.3 \( rms \) noise (1\( \sigma \)) of 6 cm images shown in Fig. 4. All five images shown in Fig. 4 have matched resolution of 9\(^\prime\) × 4\(^\prime\) at a PA of 0\(^\circ\).

<table>
<thead>
<tr>
<th>Telescope</th>
<th>Date</th>
<th>Image ( rms ) (mJy/beam)</th>
</tr>
</thead>
<tbody>
<tr>
<td>VLA</td>
<td>1984</td>
<td>0.09</td>
</tr>
<tr>
<td>VLA</td>
<td>1989</td>
<td>0.09</td>
</tr>
<tr>
<td>ATCA</td>
<td>1993</td>
<td>0.20</td>
</tr>
<tr>
<td>VLA</td>
<td>2008</td>
<td>0.06</td>
</tr>
<tr>
<td>ATCA</td>
<td>2009</td>
<td>0.06</td>
</tr>
</tbody>
</table>

### Table 4. Expansion rate of G1.9+0.3 in arc-seconds per year of radially averaged radio-continuum emission peak.

<table>
<thead>
<tr>
<th>Year</th>
<th>1984</th>
<th>1989</th>
<th>2008</th>
<th>2009</th>
</tr>
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<tr>
<td>1984</td>
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<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>1989</td>
<td>0.148</td>
<td>0</td>
<td></td>
<td></td>
</tr>
<tr>
<td>2008</td>
<td>0.189</td>
<td>0.200</td>
<td>0</td>
<td></td>
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<tr>
<td>2009</td>
<td>0.182</td>
<td>0.191</td>
<td>0</td>
<td>0</td>
</tr>
</tbody>
</table>

### Table 5. % Expansion of G1.9+0.3 per year of radially averaged radio-continuum emission peak.

<table>
<thead>
<tr>
<th>Year</th>
<th>1984</th>
<th>1989</th>
<th>2008</th>
<th>2009</th>
</tr>
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<tr>
<td>1984</td>
<td>0</td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>1989</td>
<td>0.484</td>
<td>0</td>
<td></td>
<td></td>
</tr>
<tr>
<td>2008</td>
<td>0.620</td>
<td>0.641</td>
<td>0</td>
<td></td>
</tr>
<tr>
<td>2009</td>
<td>0.598</td>
<td>0.612</td>
<td>0</td>
<td>0</td>
</tr>
</tbody>
</table>

\( ^{2} \)We combined these two frequencies into a single 1 GHz band, see the image shown in Fig. 4 (middle left).

\( ^{3} \)Distance estimates to G1.9+0.3 vary between ~7.8 kpc (Nord et al. 2004) and 10 kpc (Roy and Pal 2014). We use 8.5 kpc as the most likely distance – suggested by Reynolds et al. (2008).
Fig. 4. Matched resolution ($9'' \times 4''$ at a PA of $0^\circ$) 6-cm images of Galactic SNR G1.9+0.3 (centred at RA(J2000)=17$^h$48$^m$45.4, Dec(J2000)=−27$^\circ$10$'$06$''$) at multiple epochs. Left to right, top to bottom, 2009 ATCA, 2008 VLA, 1993 ATCA, 1989 VLA and 1984 VLA.
In Tables 4 and 5 we show the expansion of G1.9+0.3 over the period between 1984 and 2009. The mid range expansion over this period (excluding the 1993 observations) is 0.563%±0.078% yr\(^{-1}\) or 0.174±0.026 arcsec yr\(^{-1}\). However, we note that between 1984 and 1989, G1.9+0.3 expanded at a somewhat slower rate (0.484% yr\(^{-1}\)) which was also estimated by Gómez and Rodríguez (2009) at 0.46%±0.11% based on 20-cm VLA data from the same period. Since 1989, G1.9+0.3 appears to expand at a faster rate of up to 0.641% per year which is in excellent agreement with the X-ray estimates of Carlton et al. (2011a,b) of 0.642%±0.049% and Green et al. (2008) using VLA observations. This may be an indication that perhaps the shock has "broken through" to a region of lower density and thus accelerated. The mid range expansion rate derived here, yields an upper age limit of 181±25 years (Table 6) assuming a constant expansion rate since the SN event. This dates the SN event back to the year 1828 (±25).

Similarly, the speed at which the radially averaged radio-continuum emission peak is moving (assuming a distance to G1.9+0.3 of 8.5 kpc) is somewhat slower in the 1980’s (≈6,000 km sec\(^{-1}\)) than later, when it speeds up to ≈8,000 km sec\(^{-1}\) (Table 7). The speed of radially averaged radio-continuum emission peak should not be confused with the expansion velocity of the SNR shock front, which has been estimated at ≈18,000 km sec\(^{-1}\) by Borkowski et al. (2013).

As our determination of the radially averaged radio-continuum emission peak speed is dependent on the distance to the SNR we estimate that at the lower end (between 1984 and 1989) the speed varies between 5,500 km sec\(^{-1}\) (at a distance of 7.8 kpc) and 7,000 km sec\(^{-1}\) (at a distance of 10 kpc). More recently (between 1989 and 2009), the speed varies between 7,400 km sec\(^{-1}\) (at a distance of 7.8 kpc) and 9,500 km sec\(^{-1}\) (at a distance of 10 kpc).
3.2. G1.9+0.3 Spectral Energy Distribution

By matching the resolutions of our 2009 ATCA images at 6-cm and 13-cm to that of the 20-cm image (10''9 × 5''4 at PA = –0.5°), a three point spectral index map was created, allowing for the examination of the spatial spectral variations in the remnant (see Fig. 6). In this map, the colour of each pixel represents the spectral index \( \alpha \) across the three observational frequencies.

From this map we can see that the radio-continuum spectral energy distribution across the NW and SE regions is flatter \((\alpha \sim -0.5)\) which is also following the contours of the strongest of X-ray emission (so called "X-ray ears" of G1.9+0.3; see Fig. 6 – bottom right). This SED flattening in the NW and SE is also confirmed by Farnes (2012) in his VLA observations. The steeper \((\alpha \sim -1)\) radio spectra is dominant in the Northern region of G1.9+0.3 corresponding to where the radio emission is the strongest and therefore indicating the synchrotron radio-continuum emission. We also note that a somewhat steeper spectral index \((\alpha < -1.0)\) is dominating the inside part of the SNR, while flatter spectra is at the edges.

We estimate that the average spectral index across the entire SNR is \( \alpha = -0.72 \pm 0.26 \) which is flatter than LaRosa et al. (2000) \( \alpha = -0.93 \pm 0.25 \), however, their estimates are based on 20/90 cm flux density measurements. This may indicate that the synchrotron emission may be even more dominant at higher wavelengths (92 cm). Using Green et al. (2008) flux density estimates at 20 and 6 cm from their 2008 VLA observations, we estimate the spectral index \( \alpha = -0.62 \pm 0.08 \) which is in good agreement with our ATCA estimates from approximately one year later. This steeper spectral index is expected for younger SNRs (Bell et al. 2011), and further confirms its young age.
3.3. Polarisation of G1.9+0.3

Since the ATCA observations recorded Stokes parameters $Q$, $U$ and $V$, in addition to total intensity $I$, we were able to determine the polarisation of G1.9+0.3. In our 6-cm image (Fig. 7) we show the regions of polarised emission for G1.9+0.3. The electric field vectors follow the shell of the SNR around most of the circumference of the SNR, particularly along its eastern side.

The maximum fractional polarisation is estimated to be $P = 17 \pm 3\%$ with a mean of $6\%$. No reliable polarisation was detected at 20 or 13 cm. This might indicate a significant depolarisation in the remnant, however the polarimetric response of the ATCA is known to be poor at 13-cm.

![Fig. 7. 6-cm ATCA observations of G1.9+0.3. The blue ellipse in the lower-left corner represents the synthesised beam width of 10′940 × 5′384 at PA = 0°.5. The length of the vectors represents the fractional polarised intensity at each pixel position, and their orientations indicate the mean PA of the electric field (averaged over the observing bandwidth, not corrected for any Faraday rotation). The blue line below the beam ellipse represents the length of a polarisation vector of 100%. The maximum fractional polarisation is $17\% \pm 3\%$ with a mean of $6\%$. Contours at 2.2, 4.4, 6.6, 9.4, 13, 17, 21, 26 and 32 mJy beam$^{-1}$.](image)

Typically, young type Ia SNRs exhibit a radially oriented magnetic field (tangentially oriented electric field), which is to be expected from Rayleigh-Taylor instabilities in a decelerating remnant (Gull 1975, Chevalier 1976). This is consistent with similarly young Galactic SNRs, as well as in the LMC (e.g. Table 3 in Bozzetto et al. (2014)). As we have plotted electric field vectors, we would expect them to be tangential to the circumference of the remnant. We can see in Fig. 7, the orientation of the electric field vectors roughly follow this arrangement, however, it can be seen that this pattern is not strictly followed around the entire remnant. Given the location of G1.9+0.3, towards the Galactic centre, this is most likely due to Faraday depolarisation.

Farnes (2012) also detected the presence of polarised emission. Our ATCA polarimetric results also suggest that the above observed variation is most consistent with an ambient B field perpendicular to the axis of bilateral symmetry indicated by Farnes (2012). Moreover, Farnes (2012) argues that the increased ordering of the B field in the NW as the strong Faraday depolarisation must also be present.

Farnes (2012) also argues that an intrinsically radially-oriented field could be provided by a systematic gradient in Rotation Measure (RM) of 140 rad m$^{-2}$ from N to S and can also explain the depolarisation that we observe in our ATCA images.

3.4. Magnetic Field of G1.9+0.3

We used the new equipartition formulae derived by Arbutina et al. (2012, 2013) based on the diffusive shock acceleration (DSA) theory of Bell (1978) to estimate a magnetic field strength. These formulae are particularly relevant to magnetic field estimation in SNRs, and yield magnetic field strengths between those given by the classical equipartition (Pacholczyk 1970) and revised equipartition (Beck and Krause, 2005) methods. We estimate the magnetic field strength of G1.9+0.3 to be $B \approx 273 \mu G$ and the minimum total energy of the synchrotron radiation to be $E_{\text{min}} \approx 1.8 \times 10^{48}$ ergs (see Arbutina et al. 2012, 2013) and corresponding online calculator4).

For this estimate, we used a spectral index value of $\alpha = -0.72$, integrated flux density $S_{1450} = 0.935$ Jy at $\nu = 1.425$ GHz (Green et al. 2008), distance $D = 8.5$ kpc, SNR radius of $r = 46''$ and filling factor of 0.33. However, if we additionally assume a shock velocity of 18000 km s$^{-1}$ (as suggested by Borkowski et al. (2013)) than the magnetic field strength of G1.9+0.3 becomes somewhat lower ($B \approx 180 \mu G$) and the minimum total energy of the synchrotron radiation is $E_{\text{min}} \approx 7.6 \times 10^{47}$ ergs. These estimates are very similar to Arbutina et al. (2012) estimates of $B \approx 225 \mu G$ and $E_{\text{min}} \approx 9.3 \times 10^{47}$ ergs.

A large magnetic field strength such as this ($273 \mu G$) is expected for a young SNR (Bell 2004). Indeed, this makes G1.9+0.3 a remnant with one of the highest estimated magnetic field strengths known to date. For example, other young Galactic SNRs (Beck and Krause 2005, Arbutina et al. 2012 (see their Table 1)), such as Cas A ($B \approx 1250 \mu G$), Kepler ($B \approx 141 \mu G$), G349.7+0.2 ($B \approx 523 \mu G$) and Tycho ($B \approx 285 \mu G$), are known remnants with the strongest magnetic fields. Also, in a large Magellanic Cloud (LMC) SNR, J0509-6731 (also remnant from a Type Ia SN explosion) at 400 yrs age has magnetic field strength of 168 $\mu G$ (Bozzetto et al.

4http://poincare.matf.bg.ac.rs/~arbo/eqpt/
2014), while the magnetic field of LMC SNR J0519–6902 is 171 µG (Bozzetto et al. 2012). The Small Magellanic Cloud SNR HFPK 443 (Crawford et al. 2014, in press) has a field strength of 90 µG with numerous other older remnants falling below these values. It is most likely that G1.9+0.3 is going through an evolutionary stage where the magnetic field has been amplified (added to simple compression by the shocks), which may explain such a high magnetic field value (Telezhinsky et al. 2012). The amplification of magnetic field is a process driven by very fast shocks of young SNRs. Because of the strong amplification of magnetic field, a spectral index of $\alpha = -0.72$, and the location in the surface brightness-diameter diagram, this SNR is of younger age, in free expansion stage, and expanding in a low density environment.

4. CONCLUSIONS

Here, we have presented new 6-cm, 13-cm and 20-cm ATCA observations of Galactic SNR G1.9+0.3 made in 2009. Using the new 6-cm data and archival 6-cm data we observe that there are indications that the expansion of G1.9+0.3 accelerated after 1989. Our results are in a broad agreement with the estimates of expansion made by Reynolds et al. (2008), Green et al. (2008), Gómez and Rodríguez (2009) and Carlton et al. (2011a,b). We find that at ~181 yrs, G1.9+0.3 is indeed very young and most likely the youngest SNR in the Milky Way. This very young age is also reinforced by the magnetic field arrangement and strength we estimate.

We make the following findings:

- Expansion rate of 0.484% per year between 1984 and 1989;
- Expansion rate of 0.641% per year between 1989 and 2009;
- Expansion rate of 0.563%±0.078% per year between 1984 and 2009;
- Age of 181 yrs±25 yrs;
- Average spectral index between 20-cm and 6-cm, across the entire SNR is $\alpha = -0.72±0.26$;
- 6% fractionally polarised radio emission with a peak of 17%±3%;
- Magnetic field strength $B \approx 273$ µG;

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Оригинални научни рад

У овој студији представљамо нова радио-континуум посматрања телескопом ATCA на 6 cm најмлађег остатка супернових (OS) у нашој Галаксији. Анализирана су и сва доступна ATCA и VLA посматрања на овој фреквенцији. Остатак G1.9+0.3 је $\sim 181\pm 25$ година стар и годишња стопа раста му је 0.563$\pm 0.078$\% у периоду између 1984. и 1989. Приметили смо да се између 1984. и 1989. остатак ширио веће спораје (0.484\% годишње) него после 1989., када је експанзија достигла ниво од 0.641\% годишње. G1.9+0.3 има типичан радио-спектар за младе остатке са $\alpha = -0.72\pm 0.26$. G1.9+0.3 емитује просечно 6\% поляризованог зрачења на таласној дужини од 6 cm, са максималним интензитетом 17\%$\pm 3$\%. Ова поляризациона емисија тачно прати ренгенско зрачење. Процењено је прилично јако магнетно поље у вредности од 273 $\mu$G што представља једно од најснажнијих магнетних поља у до сада посматраним OS.